

THE PHYSICAL PROPERTIES OF FOUR ~ 600 K T DWARFS

S. K. LEGGETT¹, MICHAEL C. CUSHING², D. SAUMON³, M. S. MARLEY⁴, T. L. ROELLIG⁴, S. J. WARREN⁵, BEN BURNINGHAM⁶,
H. R. A. JONES⁶, J. D. KIRKPATRICK⁷, N. LODIEU⁸, P. W. LUCAS⁶, A. K. MAINZER⁹, E. L. MARTÍN^{8,10}, M. J. MCCAUGHREAN¹¹,
D. J. PINFIELD⁶, G. C. SLOAN¹², R. L. SMART¹³, M. TAMURA¹⁴, AND J. VAN CLEVE¹⁵

¹ Gemini Observatory, 670 N. A’ohoku Place, Hilo, HI 96720, USA; sleggett@gemini.edu

² Institute for Astronomy, University of Hawaii, 2680 Woodlawn Drive, Honolulu, HI 96822, USA

³ Los Alamos National Laboratory, P.O. Box 1663, MS F663, Los Alamos, NM 87545, USA

⁴ NASA Ames Research Center, Mail Stop 245-3, Moffett Field, CA 94035, USA

⁵ Imperial College London, Blackett Laboratory, Prince Consort Road, London SW7 2AZ, UK

⁶ Centre for Astrophysics Research, Science and Technology Research Institute, University of Hertfordshire, Hatfield AL10 9AB, UK

⁷ IPAC, California Institute of Technology, Mail Code 100-22, 770 South Wilson Avenue, Pasadena, CA 91125, USA

⁸ Instituto de Astrofísica de Canarias, 38200 La Laguna, Spain

⁹ Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, CA 91109, USA

¹⁰ University of Central Florida, Physics Department, PO Box 162385, Orlando, FL 32816, USA

¹¹ School of Physics, University of Exeter, Stocker Road, Exeter, Devon EX4 4QL, UK

¹² Astronomy Department, Cornell University, Ithaca, NY 14853, USA

¹³ INAF-Osservatorio Astronomico di Torino, Strada Osservatorio 20, 10025 Pino Torinese, TO, Italy

¹⁴ National Astronomical Observatory, Mitaka, Tokyo 181-8588, Japan

¹⁵ Ball Aerospace and Technologies Corp., Boulder, CO 80301, USA

Received 2008 December 9; accepted 2009 January 28; published 2009 April 7

ABSTRACT

We present *Spitzer* 7.6–14.5 μm spectra of ULAS J003402.77–005206.7 and ULAS J133553.45+113005.2, two T9 dwarfs with the latest spectral types currently known. We fit synthetic spectra and photometry to the near-through mid-infrared energy distributions of these dwarfs and that of the T8 dwarf 2MASS J09393548–2448279. We also analyze near-infrared data for another T9, CFBD J005910.82–011401.3. We find that the ratio of the mid-to near-infrared fluxes is very sensitive to effective temperature at these low temperatures, and that the 2.2 μm and 4.5 μm fluxes are sensitive to metallicity and gravity; increasing gravity has a similar effect to decreasing metallicity, and vice versa, and there is a degeneracy between these parameters. The 4.5 μm and 10 μm fluxes are also sensitive to vertical transport of gas through the atmosphere, which we find to be significant for these dwarfs. The full near- through mid-infrared spectral energy distribution allows us to constrain the effective temperature (K)/gravity (ms^{-2})/metallicity ([m/H] dex) of ULAS J0034–00 and ULAS J1335+11 to 550–600/100–300/0.0–0.3 and 500–550/100–300/0.0–0.3, respectively. These fits imply low masses and young ages for the dwarfs of 5–20 M_{Jupiter} and 0.1–2 Gyr. The fits to 2MASS J0939–24 are in good agreement with the measured distance, the observational data, and the earlier T8 near-infrared spectral type if it is a slightly metal-poor 4–10 Gyr old system consisting of a 500 K and 700 K, $\sim 25 M_{\text{Jupiter}}$ and $\sim 40 M_{\text{Jupiter}}$, pair, although it is also possible that it is an identical pair of 600 K, 30 M_{Jupiter} , dwarfs. As no mid-infrared data are available for CFBD J0059–01 its properties are less well constrained; nevertheless it appears to be a 550–600 K dwarf with $g = 300\text{--}2000 \text{ ms}^{-2}$ and [m/H] = 0–0.3 dex. These properties correspond to mass and age ranges of 10–50 M_{Jupiter} and 0.5–10 Gyr for this dwarf.

Key words: stars: atmospheres – stars: individual (ULAS J003402.77–005206.7, ULAS J133553.45+113005.2, 2MASS J09393548–2448279, CFBD J005910.82–011401.3) – stars: low-mass, brown dwarfs

1. INTRODUCTION

The low-temperature field brown dwarfs are of interest for many reasons. They are important for studies of star formation and the initial mass function (IMF) at low masses, and of the chemistry and physics of cool atmospheres. Furthermore, they populate the substellar temperature range and can act as proxies for giant exoplanets which have similar temperatures and radii, but which have less well-understood interiors and which are harder to observe than isolated brown dwarfs.

The far-red Sloan Digital Sky Survey (SDSS; York et al. 2000), and near-infrared Deep Near-Infrared Survey of the Southern Sky (DENIS; Epchtein 1998) and Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006), led to the discovery of the field population of very low mass and low temperature stars and brown dwarfs—the L and T dwarfs (e.g., Kirkpatrick et al. 1997; Martín et al. 1999; Strauss et al. 1999; Leggett et al. 2000). The L dwarfs have effective temperatures (T_{eff}) of 1400–2300 K and the T dwarfs have $T_{\text{eff}} < 1400$ K (Golimowski

et al. 2004; Vrba et al. 2004), with a lower limit not yet defined. Burrows et al. (2003) calculate that several changes will occur near $T_{\text{eff}} = 400$ K which may trigger a new spectral type, including the formation of H_2O clouds, strengthening of NH_3 , and the increasing importance of the mid-infrared flux region.

The flux from the T dwarfs is predominantly emitted in the far-red through mid-infrared, and the classification scheme is based on the strength of the near-infrared H_2O and CH_4 absorption bands (Burgasser et al. 2006). At the time of writing, only 14 T dwarfs are known with spectral type later than T7, i.e., with T_{eff} less than ~ 900 K (e.g., <http://www.DwarfArchives.org>). The near-infrared spectra of the three recently discovered T9 dwarfs (Warren et al. 2007; Burningham et al. 2008; Delorme et al. 2008a), the latest spectral type currently known, hint at an NH_3 feature at 1.5 μm and suggest that we may be at the brink of defining the next spectral class, for which the letter Y has been suggested (Kirkpatrick et al. 1999).

In this paper, we present new mid-infrared spectra for two of these T9 dwarfs, ULAS J003402.77–005206.7 and

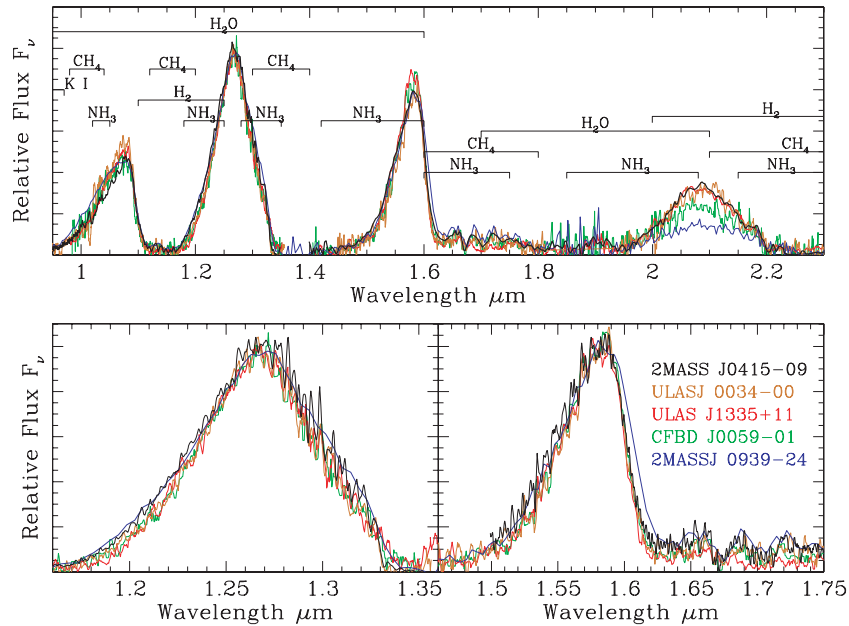


Figure 1. Near-infrared spectra of the three T9 dwarfs ULAS J0034–00, ULAS J1335+11 and CFBD J0059–01, and the T8 dwarfs 2MASS J0415–09 and 2MASS J0939–24. The top panel shows the spectra normalized to the 1.25 μm flux peak; the lower panels show the *J*- and *H*-band regions with the spectra scaled to the 1.25 μm and 1.6 μm peaks. Data sources are Burgasser et al. (2008), Burningham et al. (2008), Delorme et al. (2008a), Knapp et al. (2004), McLean et al. (2003), and Warren et al. (2007). The principal absorbing species are indicated in the top panel.

ULAS J133553.45+113005.2, obtained with the Infrared Spectrograph (IRS; Houck et al. 2004) on the *Spitzer Space Telescope* (SST; Werner et al. 2004). We combine these data with the previously published near-infrared spectra and mid-infrared Infrared Array Camera (IRAC; Fazio et al. 2004) to generate an almost complete spectral energy distribution (SED) for the dwarfs, and use atmospheric and evolutionary models to constrain their physical properties. We also investigate similar data for 2MASS J09393548–2448279, a T8 dwarf recently found to be very cool (Burgasser et al. 2008), and the near-infrared spectrum for the third T9 dwarf, CFBD J005910.82–011401.3.

2. THE SAMPLE

Two T9 dwarfs were recently found in the UK Infrared Telescope (UKIRT) Infrared Deep Sky Survey (UKIDSS; Lawrence et al. 2007) and another in the far-red Canada–France–Hawaii Telescope’s Brown Dwarf Survey (CFBDS; Delorme et al. 2008b): ULAS J003402.77–005206.7 (hereafter ULAS J0034–00; Warren et al. 2007), ULAS J133553.45+113005.2 (ULAS J1335+11; Burningham et al. 2008), and CFBD J005910.82–011401.3 (CFBD J0059–01; Delorme et al. 2008a). The two UKIDSS dwarfs were found in the 1st and 3rd data releases of the Large Area Survey (LAS) component of UKIDSS, which reaches $J = 19.6$. The 3rd data release covers 900 deg^2 of a planned 4000 deg^2 for the LAS. The CFBDS reaches $z_{\text{AB}} = 22.5$, and 350 deg^2 have been analyzed. The discovery papers show that the near-infrared spectra, as well as IRAC photometry for the two UKIDSS dwarfs, imply a very low T_{eff} of ~ 600 K for these three dwarfs. Here, we present new data for the two UKIDSS T9 dwarfs; no new data are presented for the CFBDS dwarf but we include it in our analysis as these three dwarfs appear to be similarly cool. We also include in this analysis a dwarf discovered in the 2MASS survey by Tinney et al. (2005): 2MASS J09393548–2448279 (2MASS J0939–24). Although classified as T8 using low-resolution near-infrared spectra and methane imaging, Burgasser et al.

(2008) show that mid-infrared data constrain its temperature to be similar to that of the T9 dwarfs.

Figure 1 shows the near-infrared spectra of these four dwarfs as well as the spectrum of the T8 spectral standard 2MASS J04151954–0935066 (2MASS J0415–09; Burgasser et al. 2002, 2006) which has $T_{\text{eff}} = 725\text{--}775$ K (Saumon et al. 2007). At these low temperatures very little flux remains to be absorbed in the troughs of the strong H_2O and CH_4 bands and the spectral typing based on these fluxes becomes difficult. Burningham et al. (2008), Delorme et al. (2008a), and Warren et al. (2007) show that spectral indices which measure the blue wings of the *J* and *H* flux peaks, “ W_J ” at ≈ 1.21 μm and “ $\text{NH}_3\text{--}H$ ” at ≈ 1.55 μm , appear to be useful type indicators, and that the previously defined “ $\text{H}_2\text{O--}H$ ” at ≈ 1.50 μm also retains some use at these extreme types (see, e.g., Figures 7–9 of Burningham et al.). It is likely that there is an NH_3 absorption feature in the blue wing of the *H* band for the coolest dwarfs, and Delorme et al. (2008a) suggest that these objects may be the prototype of the Y spectral class. However, the feature is difficult to distinguish from the strong H_2O and CH_4 bands and so until even later-type dwarfs are available to define the Y class we adopt a type of T9 here (as also suggested by Burningham et al.).

The three dwarfs classified as T9 have narrower *J* and *H* (1.2 and 1.6 μm) flux peaks than 2MASS J0415–09, suggesting that they are cooler than 750 K. The near-infrared indices, including the “ W_J ” and “ $\text{NH}_3\text{--}H$ ” indices, confirm an earlier type of T8 for 2MASS J0939–24. However, as shown by Burgasser et al. (2008) and as we demonstrate below, the mid-infrared data for 2MASS J0939–24 constrain its temperature to be similar to the T9 dwarfs. The near-infrared spectrum of 2MASS J0939–24 is very low resolution with $R \approx 150$ or $\Delta\lambda \approx 0.01$ μm , so that the detection of the variation in the width of the *J* and *H* bands is somewhat compromised; however, the object is further complicated by possible binarity (Burgasser et al. and Section 5.4) and we show later that multiplicity can address the discrepancy between near-infrared spectral type and temperature.

Table 1
Astrometry and Photometry of the Sample

Property	ULAS J0034–00 ^a	CFBD J0059–01	2MASS J0939–24 ^b	ULAS J1335+11
R.A.	00:34:02.77	00:59:10.90	09:39:35.48	13:35:53.45
Decl.	–00:52:06.7	–01:14:01.1	–24:48:27.9	+11:30:05.2
Epoch (yyymmdd)	20051004	20070805	20000210	20070421
μ_α (arcsec yr ^{–1})	–0.12 \pm 0.05	0.94 \pm 0.06	0.573 \pm 0.002	0.16 \pm 0.05
μ_δ (arcsec yr ^{–1})	–0.35 \pm 0.05	0.18 \pm 0.06	–1.045 \pm 0.003	–0.29 \pm 0.05
Distance (pc)	14 \pm 3		5.34 \pm 0.13	
z_{AB}	22.11 \pm 0.05	21.93 \pm 0.05		22.04 \pm 0.10
Y	18.90 \pm 0.10	18.82 \pm 0.02	16.47 \pm 0.09	18.81 \pm 0.04
J	18.15 \pm 0.03	18.06 \pm 0.03	15.61 \pm 0.09	17.90 \pm 0.01
H	18.49 \pm 0.04	18.27 \pm 0.05	15.96 \pm 0.09	18.25 \pm 0.01
K	18.48 \pm 0.05	18.63 \pm 0.05	16.83 \pm 0.09	18.28 \pm 0.03
3.55 μ m	16.28 \pm 0.03		13.76 \pm 0.04	15.96 \pm 0.03
4.49 μ m	14.49 \pm 0.03		11.66 \pm 0.04	13.91 \pm 0.03
5.73 μ m	14.82 \pm 0.05		12.96 \pm 0.04	14.34 \pm 0.05
7.87 μ m	13.91 \pm 0.06		11.89 \pm 0.04	13.37 \pm 0.07
References	1	2	3, 4	5

Notes.

^a The preliminary parallax is a private communication from R. Smart (2008).

^b The MKO-system *YJHK* photometry is synthesized from the spectrum which was flux-calibrated using 2MASS photometry.

References. (1) Warren et al. 2007; (2) Delorme et al. 2008a; (3) Tinney et al. 2005; (4) Burgasser et al. 2008; (5) Burningham et al. 2008.

Table 1 lists astrometric and photometric data for the three T9 dwarfs and 2MASS J0939–24, taken from the sources referenced in the table. The z photometry for the UKIDSS dwarfs has been converted to the SDSS AB system using transformations given by Warren et al. (2007). The *YJHK* photometry is on the Mauna Kea Observatories Vega-based system (Tokunaga et al. 2002; the *JHK* for 2MASS J0939–24 is synthesized from its spectrum which was flux calibrated using 2MASS photometry). The IRAC photometry is reported in the standard IRAC system, which is calibrated against Vega. The IRAC values provide a flux density at a nominal wavelength assuming a flux distribution where νf_ν is constant, which is not the case for T dwarfs. Hence the IRAC magnitudes do not give the flux at the nominal wavelengths listed in Table 1, but do give a scaling factor for the spectrum over the IRAC filter bandpasses (e.g., Cushing et al. 2006).

3. SPITZER IRS OBSERVATIONS

We used the Short-Low (SL) module of the IRS to obtain first-order 7.6–15.2 μ m spectra of ULAS J0034–00 and ULAS J1335+11 in the standard staring mode, which produces low-resolution spectra with $\lambda/\Delta\lambda \approx 60$ –120. Wavelengths longer than $\sim 14.5 \mu$ m are compromised by an overlapping second order. The blue pickup band was used to center the targets in the 3''7 slit by offsetting from a nearby bright star, and cycles were repeated after nodding the target along the slit by roughly one- and two-thirds of its length.

ULAS J0034–00 was observed as part of program 40419 using instrument-team guaranteed time in Cycle 4. The dwarf was observed on 2008 January 15 and January 16, each time using a 240 s ramp duration and 32 cycles. The total time on source was 8.5 hr, and the total Astronomical Observation Request (AOR) duration was 10.8 hr.

The brighter dwarf ULAS J1335+11 was observed as part of program 40449 using general observer time in Cycle 4. The dwarf was observed on 2008 August 9, using a 60 s ramp duration and 64 cycles. The total time on source was 2.1 hr, and the total AOR duration was 2.8 hr.

Low-level processing of the spectra was done automatically by versions 17.2.0 and 18.1.0 of the IRS pipeline for ULAS J0034–00 and ULAS J1335+11, respectively. The IRS pipeline removed all the instrumental artifacts to produce Basic Calibrated Data (BCD) for each exposure. The data were further reduced as described in Cushing et al. (2006). Briefly, the spectra were extracted using a modified version of the Spextool package with a fixed aperture of 6''. Observations of α Lac, obtained as part of the IRS calibration observations, were used to remove the instrument response and flux calibrate the spectra. The flux calibration was checked using the IRAC band 4 (6.4–9.3 μ m) photometry, extrapolating the observed spectrum blueward from 7.6 to 6.4 μ m using the observed spectrum of the T8 dwarf 2MASS J0415–09 as a template. Very little flux is emitted at these wavelengths due to strong absorption by H₂O at ~ 5 –7 μ m and CH₄ at ~ 7 –9 μ m. This photometric check showed that the IRS calibration is accurate to $\leq 10\%$.

Figure 2 shows the IRS spectra of ULAS J0034–00 and ULAS J1335+11.

4. ATMOSPHERIC AND EVOLUTIONARY MODELS

Atmospheric and evolutionary models were generated by members of our team (Marley et al. 2002; Saumon & Marley 2008). The atmospheric models include all the significant sources of gas opacity (Freedman et al. 2008). However, there are known deficiencies in the molecular opacity line lists at $0.9 < \lambda \mu\text{m} < 1.4$, where the line list for CH₄ is incomplete and the line list for NH₃ does not exist. At the temperatures considered here these opacities are significant and we return to this issue later in Sections 5 and 6. The models assume solar system abundances and calculate the equilibrium abundances of the important C-, N-, and O-bearing gases as detailed by Freedman et al. (2008). These equilibrium values are then used to determine the abundances of the various gas species in each atmospheric layer during the construction of the temperature–pressure profile. We iteratively compute models in radiative–convective equilibrium, solving for a temperature–pressure profile while employing the source-function technique

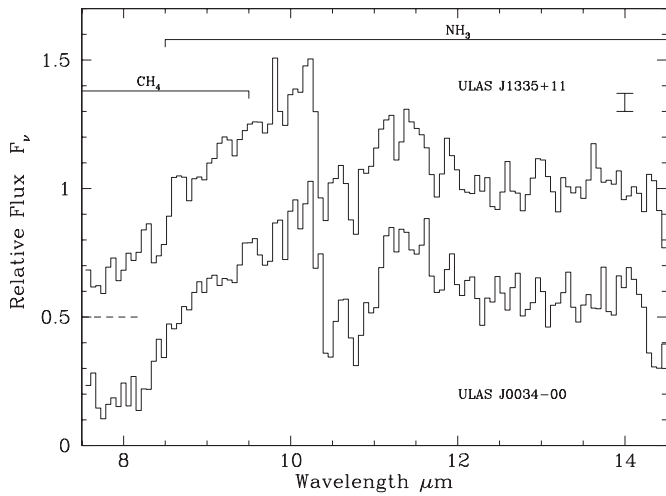


Figure 2. *Spitzer* IRS spectra of ULAS J0034–00 and ULAS J1335+11, normalized to the flux peak at $10.2\,\mu\text{m}$ and offset for clarity; a typical error bar is shown. The dashed lines show the zero flux level for ULAS J1335+11. The principal absorbing species are indicated.

of Toon et al. (1989) to solve the equation of transfer. During this process, condensates of certain refractory elements are included in atmospheric layers where physical conditions favor such condensation. At the temperatures considered here the condensate clouds are located deep below the photosphere and the opacity of the clouds is ignored although grain condensation is still accounted for in the chemistry.

The models also include vertical transport in the atmosphere, which significantly affects the chemical abundances of C-, N-, and O-bearing species. The very stable molecules CO and N_2 are dredged up from deep layers to the photosphere, causing enhanced abundances of these species and decreased abundances of CH_4 , H_2O , and NH_3 (e.g., Fegley & Lodders 1996; Hubeny & Burrows 2007; Saumon et al. 2007). Vertical transport is parameterized in our models by an eddy diffusion coefficient K_{zz} ($\text{cm}^2\,\text{s}^{-1}$) that is related to the mixing timescale. Generally, larger values of K_{zz} imply greater enhancement of CO and N_2 over CH_4 and NH_3 , respectively. Values of $\log K_{zz} = 2$ –6, corresponding to mixing timescales of ~ 10 yr to ~ 1 hr, respectively, reproduce the observations of T dwarfs (e.g., Leggett et al. 2007b; Saumon et al. 2007).

The evolutionary sequences of Saumon & Marley (2008), computed with cloudless models as the surface-boundary condition, provide radius, mass, and age for a variety of atmospheric parameters. For this investigation, synthetic spectra were generated with T_{eff} of 500, 550, 600, 650, and 700 K; gravity g of 100, 300, 1000, and 2000 $\text{m}\,\text{s}^{-2}$ (corresponding to $\log g$ in cgs units of 4.00, 4.48, 5.00, and 5.30); metallicities of solar and twice solar ($[\text{m}/\text{H}] = +0.3$); and vertical mixing diffusion coefficient K_{zz} of zero (corresponding to equilibrium) and $10^4\,\text{cm}^2\,\text{s}^{-1}$. A limited set was also calculated with $K_{zz} = 10^6\,\text{cm}^2\,\text{s}^{-1}$ and $[\text{m}/\text{H}] = -0.3$. The evolutionary models show that a brown dwarf with $T_{\text{eff}} \leq 700$ K must have $g \leq 2000\,\text{m}\,\text{s}^{-2}$ if its age is ≤ 10 Gyr (Saumon & Marley 2008).

Figure 3 plots a selection of 500 K and 600 K synthetic spectra to demonstrate the effects of varying the atmospheric parameters. The top panel shows that metallicity and gravity variations can affect the SEDs similarly—an increase in metallicity or a decrease in gravity leads to significantly increased flux in the K band at $2.2\,\mu\text{m}$, and decreased flux at $4.5\,\mu\text{m}$ (impacting IRAC band 2 photometry for example). The opposite changes occur if

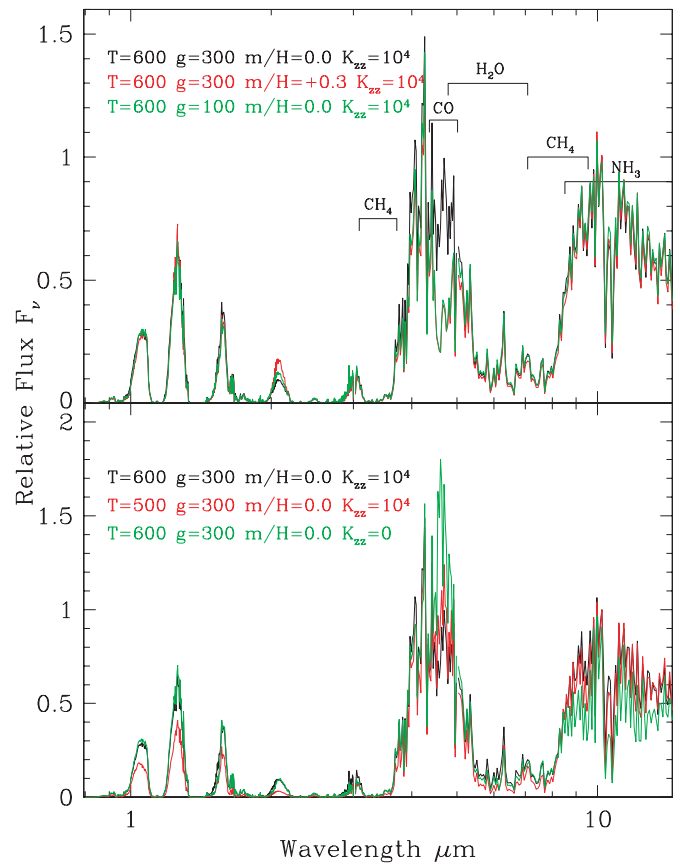


Figure 3. Synthetic spectra demonstrating the effect of varying the atmospheric parameters; the spectra are normalized to unity at $\lambda = 10\,\mu\text{m}$. In both panels, the black curves are the baseline model with $T_{\text{eff}} = 600$ K, $g = 300\,\text{m}\,\text{s}^{-2}$, $[\text{m}/\text{H}] = 0$, and $K_{zz} = 10^4\,\text{cm}^2\,\text{s}^{-1}$. The top panel demonstrates the similar effects of increasing metallicity and decreasing gravity—the $2.2\,\mu\text{m}$ flux is increased while the $4.5\,\mu\text{m}$ flux is decreased. The lower panel demonstrates that reducing T_{eff} reddens the slope of the energy distribution, and removing chemical mixing increases the $4.5\,\mu\text{m}$ flux while decreasing the 9 – $15\,\mu\text{m}$ flux. The principal absorbing species in the mid-infrared are indicated in the top panel.

metallicity is decreased or gravity is increased. The lower panel shows that varying T_{eff} impacts the overall slope of the SED—decreasing T_{eff} increases the ratio of the mid- to near-infrared fluxes (although surface flux at all wavelengths decreases with decreasing temperature). The lower panel also shows that the mixing parameter K_{zz} strongly affects the $4.5\,\mu\text{m}$ and 9 – $15\,\mu\text{m}$ fluxes: increasing K_{zz} reduces the $4.5\,\mu\text{m}$ flux as CO is dredged up, and increases the 9 – $15\,\mu\text{m}$ flux as N_2 is favored over NH_3 . For the objects considered here, the CH_4 and H_2O bands are not significantly affected by CO mixing, as the amount of carbon and oxygen tied up in CO is a small fraction of that in these other molecules at these temperatures.

5. ANALYSIS

5.1. The 0.9–2.5 μm and 7.6–14.5 μm Spectra of ULAS J0034–00 and ULAS J1335+11

The synthetic spectra were compared to the observed 0.9–2.5 μm and 7.6–14.5 μm spectra of ULAS J0034–00 and ULAS J1335+11 using the fitting procedure described by Cushing et al. (2008), except that here we do not mask out the 1.58 to $1.75\,\mu\text{m}$ spectral region. No regions were ignored in the fitting, despite the known problems at $0.9 < \lambda\,\mu\text{m} < 1.4$ (Section 4). The procedure uses weighted least-squares fitting to identify the synthetic spectrum that best matches the observed

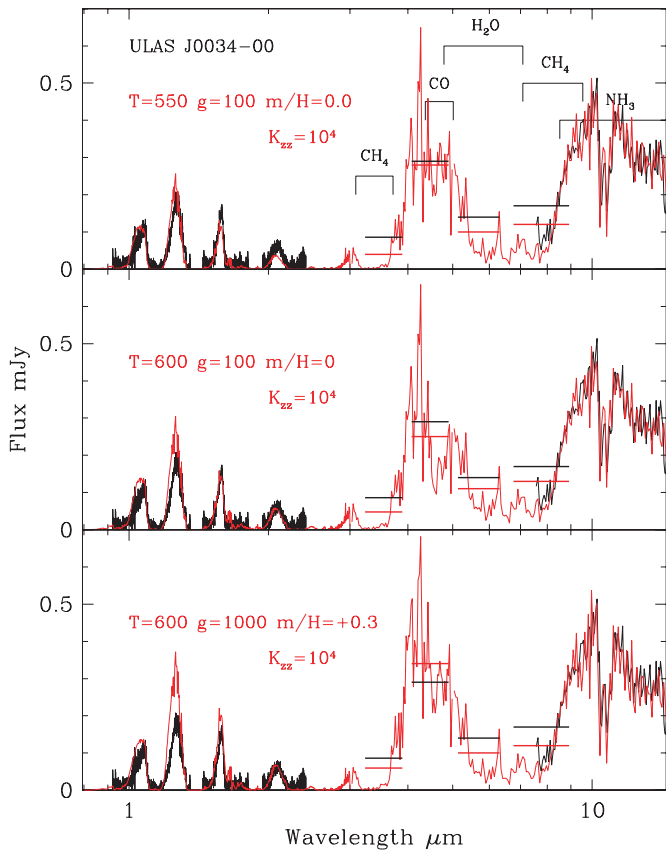


Figure 4. Near- and mid-infrared spectra of ULAS J0034–00 (black lines) compared to various synthetic spectra (red lines) generated by models with parameters as indicated in the legends. Scaling of the models to the data is done by minimizing the weighted least-squares fit; this provides a distance to the dwarf. The horizontal lines show the observed and calculated IRAC fluxes at this distance. The uncertainty in the measured IRAC fluxes is 3%–6%. A gravity as high as $g = 1000 \text{ m s}^{-2}$ is excluded by the poorer spectral fit shown in the bottom panel.

spectrum. The scaling of the synthetic spectra to the observations is determined as part of the fitting procedure, by minimizing the weighted deviation between the two data sets.

Fits using the entire wavelength range as well as only the near- or mid-infrared were carried out. For both of these T9 dwarfs the fits to either the near- or mid-infrared regions alone imply a value of T_{eff} of $600 \leq T_{\text{eff}} \leq 700$. However, the fit to the mid-infrared spectrum does not constrain the temperature strongly, and the fit to the near-infrared is very poor in the H band. The fits to the full spectrum give a lower temperature of $500 \leq T_{\text{eff}} \leq 600$ for ULAS J0034–00 and $500 \leq T_{\text{eff}} \leq 550$ for ULAS J1335+11; in these full-range fits the match to the H band is improved at the expense of a poorer fit to the K band. Increasing T_{eff} makes the overall spectrum bluer (e.g., Figure 3), and temperatures as high as suggested by the near-infrared fits cannot reproduce the observed full SEDs.

The best fitting model to the spectrum of ULAS J0034–00 has effective temperature (K)/gravity (m s^{-2})/metallicity ([m/H] dex) of 550/100/0.0; this fit is shown in the top panel of Figure 4. However, temperatures in the range 550–600 K, gravities of 100–300 m s^{-2} and metallicities of 0.0–0.3 dex do not give significantly worse fits, as shown in the top two panels of Figures 4 and 5. The best fitting model to the spectrum of ULAS J1335+11 has effective temperature (K)/gravity (m s^{-2})/metallicity ([m/H] dex) of 500/100/0.0; this fit is shown in the top panel of Figure 6. Temperatures in the range 500–550 K,

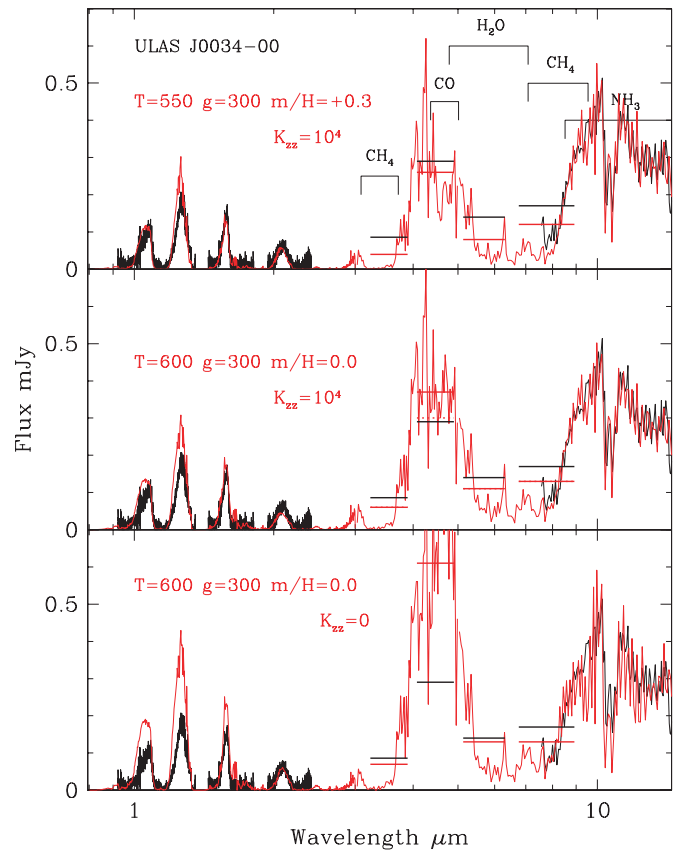


Figure 5. Near- and mid-infrared spectra of ULAS J0034–00 (black lines) compared to various synthetic spectra and photometry (red lines) as in Figure 4. The dotted red lines in the middle panel shows the effect on the IRAC fluxes of increasing the mixing coefficient to $K_{zz} = 10^6 \text{ cm}^2 \text{ s}^{-1}$. Equilibrium chemistry, $K_{zz} = 0 \text{ cm}^2 \text{ s}^{-1}$, models can be excluded; the bottom panel shows that in this case reproducing the mid-infrared spectrum leads to too high near-infrared flux levels, and also a discrepancy of a factor of ~ 2 at $4.5 \mu\text{m}$ (see also the lower panel of Figure 3).

gravities of 100–300 m s^{-2} , and metallicities of 0.0–0.3 dex give similar quality fits as shown in Figure 6.

The full spectral range supports a low gravity for both dwarfs of $100 \leq g \text{ m s}^{-2} \leq 300$. Although an increase in gravity can be partly compensated for by increasing metallicity, a gravity as high as $g = 1000 \text{ m s}^{-2}$ gives a poorer least-squares fit to both dwarfs, and this can be seen for ULAS J0034–00 in the bottom panel of Figure 4.

For both dwarfs, models which include gas transport, i.e., nonequilibrium chemistry with $K_{zz} > 0$, are strongly preferred by the fits to the mid-infrared only and the fits to the entire spectral range. The equilibrium chemistry models do not reproduce the shape of the 9–15 μm spectrum, where NH_3 absorption is strong, but not as strong as it would be in the $K_{zz} = 0$ case; this is shown in the bottom panel of Figure 5. The flux at $4.5 \mu\text{m}$ is also dependent on the value of K_{zz} , which we discuss further when considering the IRAC photometry below. Vertical mixing in ULAS J1335+11 appears stronger than in ULAS J0034–00 based on the $4.5 \mu\text{m}$ IRAC photometry; we determine $K_{zz} \approx 10^6 \text{ cm}^2 \text{ s}^{-1}$ compared to $10^4 \text{ cm}^2 \text{ s}^{-1}$ for ULAS J0034–00 (Section 5.2).

5.2. The 3–9 μm Photometry of ULAS J0034–00 and ULAS J1335+11

The IRAC photometry for ULAS J0034–00 and ULAS J1335+11 (Warren et al. 2007; Burningham et al. 2008) fills the gap

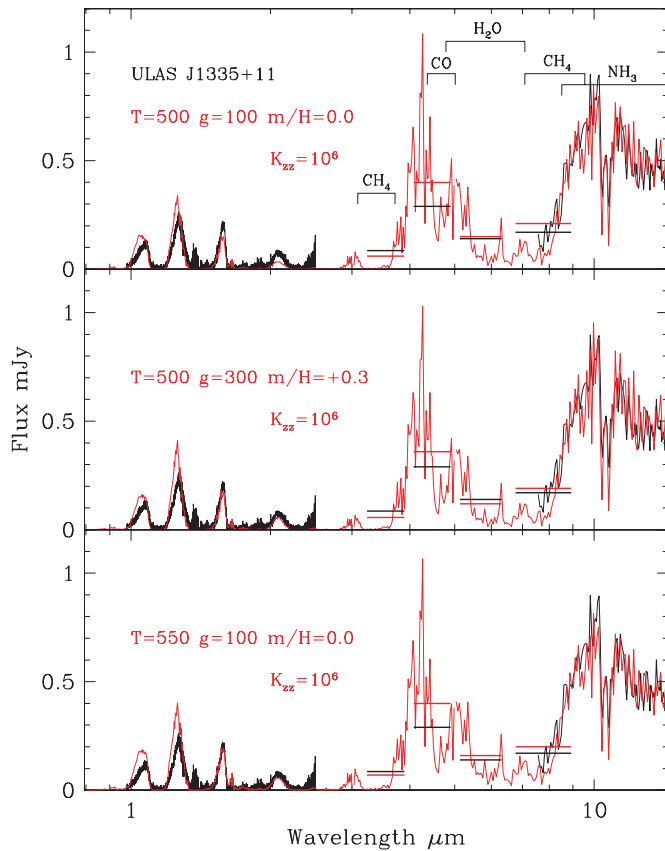


Figure 6. Fits (red lines) to ULAS J1335+11 (black lines) using models with parameters as indicated in the legends, as in Figure 4. The uncertainty in the measured IRAC fluxes is 3%–7%.

between the near- and mid-infrared spectra and provides another check of the model fit. The observed and calculated IRAC fluxes are indicated in Figures 4–6. The observed fluxes are determined by transforming the observed magnitudes to fluxes using the Vega fluxes of 280.9, 179.7, 115.0, and 64.1 Jy at the four IRAC bands [3.6], [4.5], [5.8], and [8.0], respectively, as provided in the IRAC Data Handbook. The calculated flux is determined in the same way from synthetic absolute magnitudes, scaling by the distance to the dwarf, D , which is provided by the spectral fit. The scaling factor required to match the synthetic spectra to the data is a function of the solid angle (R/D) where R is the radius of the dwarf. As R is known from the evolutionary model corresponding to the particular set of atmospheric parameters T_{eff} , g , and $[m/H]$, the distance D can be derived. Both dwarfs are assumed to be single in this analysis; both have been imaged at high resolution and not resolved (N. Lodieu 2007, private communication; M. Liu 2008, private communication).

The IRAC [4.5] flux is sensitive to the vertical transport coefficient K_{zz} as shown in Figures 3 and 5. Increasing $\log K_{zz}$ from 4 to 6 reduces the [4.5] flux by 10%–25% at these temperatures and gravities, as the nonequilibrium abundance of CO increases. The other IRAC bands are insensitive to K_{zz} at these low temperatures, as described in Section 4 and shown in the middle panel of Figure 5. Hence $\sim 20\%$ effects at [4.5] can be neglected when examining the model fits as these can be addressed by varying K_{zz} .

All the models preferred by the spectral fits to ULAS J0034–00 give IRAC [3.6], [5.8], and [8.0] fluxes which are too low by 30%–50%. The agreement at [4.5] is good as long as $K_{zz} > 0$ (which is required by the spectral fits also). The fits to

the [3.6], [5.8], and [8.0] fluxes are better for ULAS J1335+11, while the [4.5] model flux is high by 20%–30% for this dwarf, even with $K_{zz} = 10^6 \text{ cm}^2 \text{ s}^{-1}$. The fact that the subsets of the modeled and observed IRAC photometry agree well for each dwarf suggests that a more detailed model grid and a model-fitting routine which incorporates both the spectroscopy and photometry could provide a more consistent set of observed and synthetic data. We postpone such work until the molecular line lists are improved (see Section 6).

Despite the remaining $\sim 30\%$ discrepancies at [3.6], [5.8], and [8.0] for ULAS J0034–00 (regions where there is very little flux), the near- and mid-infrared spectra, and the IRAC photometry, together with our models, can satisfactorily constrain T_{eff} (K)/ g (m s^{-2})/ $[m/H]$ (dex) to 550–600/100–300/0.0–0.3 and 500–550/100–300/0.0–0.3 for ULAS J0034–00 and ULAS J1335+11, respectively. Vertical transport of gas must be significant, with a diffusion coefficient of $K_{zz} = 10^4$ – $10^6 \text{ cm}^2 \text{ s}^{-1}$, similar to values found for earlier-type T dwarfs.

These parameters and their corresponding radii, masses, ages, distances, and tangential velocities are given in Table 2. The low temperatures and gravities imply that the dwarfs have masses of only 5–20 M_{Jupiter} and young ages of 0.1–2 Gyr. We discuss the uncertainties in these parameters in Section 6. We note that the implied distance of 13–16 pc for ULAS J0034–00 is in excellent agreement with a preliminary trigonometric parallax of 14 ± 3 pc (R. Smart 2008, private communication).

5.3. Trends in Color

Fitting the synthetic spectra to the observations showed that the ratio of the mid- to near-infrared flux is very sensitive to temperature. Combining ground-based near-infrared photometry with IRAC photometry therefore should be a useful indicator of T_{eff} , in particular the [4.5] band as there is very little flux at [3.6], [5.8], and [8.0]. Warren et al. (2007) investigate various color combinations and find that H –[4.5] is a good indicator of T_{eff} for T dwarfs. While the K band is sensitive to metallicity and gravity, the H band is insensitive (e.g., Figure 3). Stephens et al. (2009) show that T dwarf colors involving J and IRAC magnitudes show a lot of scatter with T_{eff} , at least some of which is due to condensates clearing out of the photospheric region probed by $1.2 \mu\text{m}$ radiation in the earlier T dwarfs.

Figure 7 shows H –[4.5] as a function of H – K . Model sequences are shown for a range of gravity and metallicity, and for nonequilibrium chemistry models with $K_{zz} = 10^4 \text{ cm}^2 \text{ s}^{-1}$. Increasing K_{zz} by a factor of 100 leads to no significant change in H – K , but a decrease in H –[4.5] of 0.1–0.3 mag. From this and other studies it appears that the range of K_{zz} in T dwarfs is 10^2 – $10^6 \text{ cm}^2 \text{ s}^{-1}$ and so the uncertainty in the modeled H –[4.5] due to unknown K_{zz} should be ≈ 0.2 mag. Plausible changes in gravity and metallicity for a local disk sample also affect H –[4.5] by ~ 0.2 mag and affect H – K by ~ 0.3 – 0.6 mag at the temperatures considered here (note that the effect on H – K increases with decreasing temperature). As described before, the trends with metallicity and gravity are degenerate. Although these ~ 0.2 mag uncertainties are not ideal in a temperature indicator, Figure 7 shows that for $T_{\text{eff}} < 700 \text{ K}$ H –[4.5] increases by 0.6–0.8 mag per 100 K and so temperatures can still be accurately determined. The *Spitzer* warm mission retains IRAC [4.5] capability and hence will continue to provide important data for studies of brown dwarfs.

Colors for ULAS J0034–00, ULAS J1335+11, and 2MASS J0939–24 are shown in Figure 7, as well as those of other

Table 2
Likely Range of Properties of the Sample

Name	T_{eff} (K)	g (m s^{-2})	[m/H]	R/R_{\odot}	Mass (Jupiter)	Age (Gyr)	Distance (pc)	V_{tan} (km s^{-1})
ULAS J0034–00	550	100	0.0	0.12	5–8	0.1–0.2	14	25
	550	300	+0.3	0.11	10–20	1–2	13	23
	600	100	0.0	0.12	5–8	0.1–0.2	16	28
	600	300	0.0	0.11	10–20	1–2	13	25
ULAS J1335+11	500	100	0.0	0.12	5–8	0.1–0.2	9	17
	500	300	+0.3	0.11	10–20	1–2	9	14
	550	100	0.0	0.12	5–8	0.1–0.2	11	17
	550	300	+0.3	0.11	10–20	1–2	9	14
2MASS J0939–24AB ^a	600	1000	–0.3–0.0	0.09	20–40	2–10	5.34	30
2MASS J0939–24A ^b	700	1000–2000	–0.3–0.0	0.08–0.09	30–50	4–10	5.34	30
2MASS J0939–24B ^b	500	1000	–0.3–0.0	0.09	20–30	4–10	5.34	30
CFBD 0059–01	550	300	0.0	0.11	10–20	0.5–2	11	50
	550	1000	+0.3	0.09	25–40	4–10	11	50
	550	2000	+0.3	0.08	40–50	10	10	45
	600	300	0.0	0.11	10–20	0.5–2	14	64
	600	1000	0.0	0.09	25–40	4–10	11	50
	600	2000	0.0	0.08	40–50	10	10	45

Notes.

^a Assuming that this object is a pair of similar luminosity dwarfs.

^b Assuming that this object is a pair of dissimilar dwarfs.

dwarfs taken from the literature (Patten et al. 2006; Leggett et al. 2007b; Luhman et al. 2007). The T7.5 dwarfs Gl 570D and HD 3651B are benchmark dwarfs with well constrained ages and metallicities from their primary stars, and well constrained T_{eff} based on their measured luminosity and the known distance to the primary star (Burgasser et al. 2000; Burgasser 2007; Liu et al. 2007; Mugrauer et al. 2006; Saumon et al. 2006). Gl 570D has $T_{\text{eff}} = 750\text{--}825$ K, $g = 630\text{--}2000$ m s^{-2} , and $[\text{m}/\text{H}] = 0.01\text{--}0.10$; HD 3651B has $T_{\text{eff}} = 780\text{--}840$ K, $g = 1260\text{--}3160$ m s^{-2} , and $[\text{m}/\text{H}] = 0.09\text{--}0.16$. 2MASS J0415–09 has a well-measured distance and SED which indicates $T_{\text{eff}} = 725\text{--}775$ K, $g = 1000\text{--}2500$ m s^{-2} , and $[\text{m}/\text{H}] = 0.0\text{--}0.3$ dex (Saumon et al. 2007). If the sequences shown in Figure 7 are shifted redward in $H - K$ by ~ 0.2 mag and redward in $H - [4.5]$ by ~ 0.1 mag there is good agreement with the observations of these slightly metal-rich dwarfs, and with the overall population in the diagram. The models appear to be too faint at K , and this can also be seen in Figures 4–6. Most likely this is due to some inadequacy in the CH_4 line list or the H_2 collision-induced absorption, as discussed in Saumon & Marley (2008; their Section 4.1).

The color trends shown in Figure 7 support the temperatures, gravities, and metallicities found above for ULAS J0034–00 and ULAS J1335+11: that they are relatively low-gravity dwarf with solar or enhanced metallicities, and that ULAS J1335+11 is cooler than ULAS J0034–00. The colors also support the conclusion of Burgasser et al. (2008) in their analysis of the T8 dwarf 2MASS J0939–24: that it has a similar temperature to these T9 dwarfs, but a higher gravity and/or lower metallicity. We explore fits to this object’s spectral distribution in the following section.

5.4. Fitting the SED of 2MASS J0939–24

Burgasser et al. (2008) find the following parameters for 2MASS J0939–24 using near- through mid-infrared observational data: $T_{\text{eff}} = 600 \pm 50$ K, $g = 600$ m s^{-2} , and $[\text{m}/\text{H}] = -0.4$ (the uncertainties in g and $[\text{m}/\text{H}]$ are quite large). The same model set and a similar least-squares fitting routine were used as we use here for ULAS J0034–00 and ULAS J1335+11. In this case, however, the distance to the dwarf is known and so

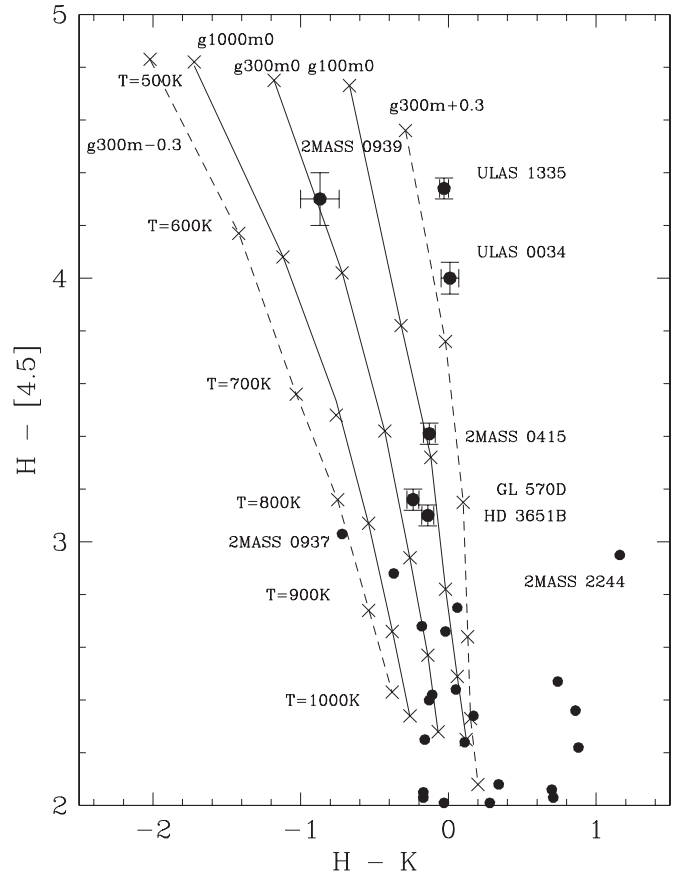


Figure 7. Calculated sequences with T_{eff} in the colors $H - K$: $H - [4.5]$ are shown as solid and dashed lines. The solid lines are solar metallicity models with gravities 1000, 300, and 100 m s^{-2} from left to right. The dashed lines have $g = 300$ m s^{-2} and demonstrate the effect of varying metallicity: $[\text{m}/\text{H}] = -0.3$ on the left and $[\text{m}/\text{H}] = +0.3$ on the right. The crosses mark $T_{\text{eff}} = 500, 600, 700, 800, 900$, and 1000 K from top to bottom. All sequences are for nonequilibrium chemistry with $K_{zz} = 10^4$ $\text{cm}^2 \text{s}^{-1}$. At these temperatures the range of gravities shown corresponds to an age range of $\sim 0.1\text{--}5$ Gyr (right to left). The solid points are the observed colors for the dwarfs discussed in the text and others taken from the literature. See Section 5.3.

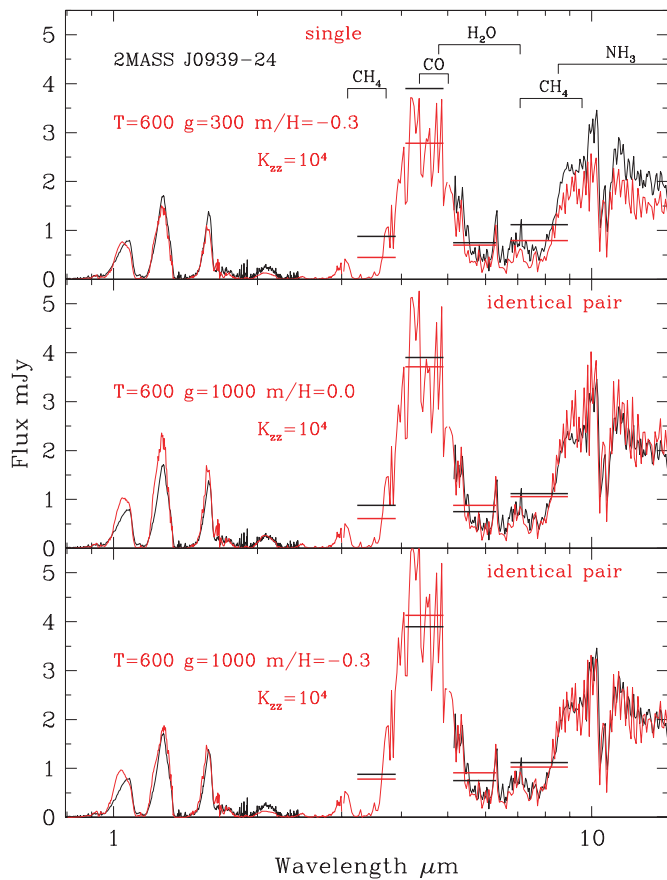


Figure 8. Fits (red lines) to 2MASS J0939–24 (black lines) using models with parameters as indicated in the legends. The top panel assumes that the dwarf is a single object and the lower panels assume that it is an identical pair of objects. The horizontal lines show the observed and calculated IRAC fluxes. The uncertainty in the measured IRAC fluxes is 4%.

the scaling factor required to match the models to the observations gives the radius of the dwarf, or alternatively if the radius is known from evolutionary arguments then the scaling factor can be derived. Adopting the first approach Burgasser et al. (2008) show that the radius is larger than expected from the models, and they suggest that the dwarf is a binary.

Figure 8 shows this dwarf’s observed spectroscopy, IRAC photometry, and a selection of the modeled data. We have selected models with $T_{\text{eff}} = 600$ K and values of gravity and metallicity within the range determined by the Burgasser et al. (2008) least-squares fitting. Scaling of the modeled spectra and photometry is achieved using the known distance to the dwarf and the radius from evolutionary models. We assume that either it is single or an equal luminosity unresolved pair, as best matches the data. The lower gravity fit, which would be marginally consistent with this object being single, is a significantly poorer fit to the spectra and photometry than the higher gravity fits, which require it to be a binary system. Also a warmer model with $T_{\text{eff}} = 650$ K can reproduce the near-infrared data well if the dwarf is single, but is underluminous in the mid-infrared by $\sim 30\%$. A reasonable fit can be found if indeed, as suggested by Burgasser et al., this is a pair of similar luminosity 600 K dwarfs with $g \approx 1000 \text{ m s}^{-2}$ and $[m/H] = -0.3$ – 0.0 dex.

Given that a temperature as low as 600 K is unexpected for an object with the same near-infrared spectral type as the 750 K dwarf 2MASS J0415–09, and that the near-infrared data

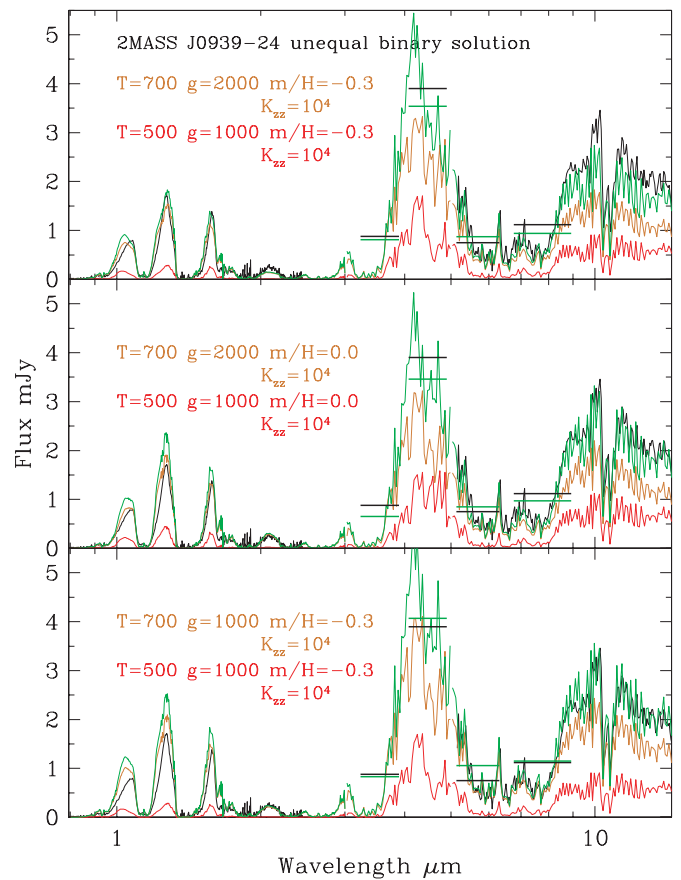


Figure 9. Fits to 2MASS J0939–24 (black line) as a nonequal luminosity binary. The 500 K component is shown in red, the 700 K in brown, and the combined spectrum in green. The horizontal lines show the observed and calculated IRAC fluxes where the IRAC fluxes are synthesized from the combined model spectrum. The uncertainty in the measured IRAC fluxes is 4%.

led Leggett et al. (2007a) to assign $T_{\text{eff}} = 750$ K to 2MASS J0939–24, we explored the possibility that this object consists of a more dissimilar pair of dwarfs. We find that a satisfactory fit is produced if 2MASS J0939–24 is composed of a 700 K and a 500 K dwarf. The gravities must be high so that the radii are small, otherwise the total flux of the system is too large. The relative gravities are also constrained by the evolutionary models, as the dwarfs presumably are coeval. Fits with $g = 1000 \text{ m s}^{-2}$ for the 500 K dwarf constrains the gravity of the 700 K dwarf to 1000 – 2000 m s^{-2} . The system still appears to be somewhat metal-poor with $[m/H] = -0.3$ – 0.0 dex. Figure 9 shows composite fits for this range of parameters. To improve the agreement with the previously determined temperature of this object, and to make the temperature:spectral type correlation stronger, this dissimilar-pair solution is preferred to the equal-luminosity pair solution. The near-infrared spectrum is then dominated by the warmer dwarf (see Figure 9) and becomes consistent with the T8 spectral type. The modeled dissimilar pair has slightly wider J and H flux peaks than the equal luminosity pair, consistent with the wider peaks observed for 2MASS J0939–24 and 2MASS J0415–09 compared to the three T9 dwarfs (Figure 1).

However, the detailed spectral shape is sensitive to gravity and metallicity as well as temperature, and the temperature:spectral type relationship will have some scatter. Hence unless the object is resolved it cannot be determined whether the equal or nonequal pair solution is correct. The physical properties of

both solutions are listed in Table 2. The 500 K + 700 K pair are constrained to be high gravity and therefore quite old at 4–10 Gyr; perhaps appropriate for the slightly metal-poor nature of the system.

5.5. The 0.9–2.5 μm Spectrum of CFBD J0059–01

At this time there are no mid-infrared data published for the T9 dwarf CFBD J0059–01. Figure 1 shows that the near-infrared spectrum of CFBD J0059–01 is very similar to ULAS J0034–00 and ULAS J1335+11, except that the K -band flux is suppressed. Burningham et al. (2008) show that the “ W_J ,” “ $\text{NH}_3\text{-}H$,” and “ $\text{H}_2\text{O-H}$ ” indices place CFBD J0059–01 consistently between ULAS J0034–00 and ULAS J1335+11, and Delorme et al. (2008a) conclude that it is ~ 50 K cooler than ULAS J0034–00 based on a comparison of near-infrared features. The indices and K flux suggest that CFBD J0059–01 has a temperature between that of ULAS J0034–00 and ULAS J1335+11 and either a higher gravity or lower metallicity. The $H - K$ value is intermediate between the ULAS T9 dwarfs and 2MASS J0939–24 (see Figure 6), suggesting that either it is not as metal-poor or as high gravity as the latter dwarf (or pair of dwarfs). Delorme et al. (2008a) find that CFBD J0059–01 is unresolved at high resolution and so here we assume that it is single.

Fitting the near-infrared spectrum alone, we find that a temperature as low as 500 K does not give a satisfactory fit; $T_{\text{eff}} = 550\text{--}600$ K produces similar fits in the near-infrared as the other dwarfs, when $g = 300\text{--}2000 \text{ m s}^{-2}$ and $[\text{m}/\text{H}] = 0\text{--}0.3$ dex. The high metallicities are required if $T_{\text{eff}} = 550$ K and $g = 1000\text{--}2000 \text{ m s}^{-2}$, to fit the K -band flux. This temperature range of 550–600 K is in agreement with the trends in the near-infrared indices, and Delorme et al.’s conclusion that CFBD J0059–01 is cooler than ULAS J0034–00, if CFBD J0059–01 is at the low end of the range while ULAS J0034–00 is at the high end.

The range of implied distances and other physical properties for CFBD J0059–01 are given in Table 2. This dwarf may be significantly older and more massive than ULAS J0034–00 and ULAS J1335+11; its tangential velocity is also significantly higher ($\sim 50 \text{ km s}^{-1}$ compared to 25 km s^{-1}). Mid-infrared data will greatly help constrain its parameters. Such data are needed to confirm the temperature implications of the near-infrared indices, and will also constrain metallicity and gravity by fixing the scaling of the models and restricting the slope of the near-infrared flux distribution.

6. DISCUSSION AND CONCLUSIONS

We have used new and published near- through mid-infrared spectroscopy and photometry to constrain the atmospheric parameters of the very low-temperature dwarfs ULAS J0034–00, ULAS J1335+11, and 2MASS J0939–24. We find strong evidence that photospheric mixing continues to be an important process, even at these cool temperatures. Any $5 \mu\text{m}$ survey for Y dwarfs (e.g., WISE or *Spitzer* warm missions) must take this mixing into account, as the flux will be less than predicted by equilibrium chemistry. Considering the complexity of the atmospheres, the models reproduce the data well. However, there are known deficiencies in the molecular opacity line lists, and discrepancies can be seen in Figures 4–6 and 8. The known problems are large at $0.9 < \lambda \mu\text{m} < 1.4$, where the line list for CH_4 is incomplete, the line list for NH_3 does not exist, and where the treatment of the red wing of the very broad $0.77 \mu\text{m K I}$ line

is difficult. These problems most likely account for the excess model flux seen in our fits at $1.0 < \lambda \mu\text{m} < 1.3$.

Given these deficiencies, the systematic errors in our derived parameters are difficult to determine. However, the good agreement with evolutionary models for the ~ 800 K brown dwarf companions to stars of known age and metallicity means that it is unlikely that the errors are larger than 100 K in T_{eff} , or a factor of 2 in gravity and metallicity (see the trends shown in Figures 3 and 7). Nevertheless, it is of paramount importance that the line lists for CH_4 and NH_3 are improved, now that objects as cool as 500 K are being discovered.

It is also very helpful at these low temperatures to have mid-infrared observational data. The absorption features in the near-infrared are effectively saturated, and the changes with temperature have become subtle. The mid-infrared flux level is very sensitive to temperature and can remove this degeneracy; this sensitivity increases with decreasing temperature as shown in Figure 7. Mid-infrared data offers the possibility of identifying low-temperature systems such as 2MASS J0939–24, which may be a 500 K + 700 K brown dwarf pair. We find that two of the three currently known T9 dwarfs have $550 \leq T_{\text{eff}} \text{ K} \leq 600$ and the third appears to be even cooler with $500 \leq T_{\text{eff}} \text{ K} \leq 550$. Vertical transport of gas remains significant even at these very low temperatures, so that enhanced CO opacity reduces the flux at $4.5 \mu\text{m}$ and decreased NH_3 opacity increases the flux at $9\text{--}15 \mu\text{m}$.

The evolutionary models show that the parameters we derive for ULAS J0034–00 and ULAS J1335+11 imply that they are both very low mass and relatively young—their masses are between 5 and $20 M_{\text{Jupiter}}$ corresponding to an age range of 0.1–2.0 Gyr. The ranges for CFBD J0059–01 are larger—10–50 M_{Jupiter} and 0.5–10 Gyr. If 2MASS J0939–24 consists of a 500 K + 700 K binary then it is relatively old and massive: a 4–10 Gyr system composed of a $25 M_{\text{Jupiter}}$ and a $40 M_{\text{Jupiter}}$ brown dwarf pair.

This work is based on observations made with the *Spitzer Space Telescope*, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. Support for this work was provided by NASA through an award issued by JPL/Caltech. S.K.L.’s research is supported by the Gemini Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., on behalf of the international Gemini partnership of Argentina, Australia, Brazil, Canada, Chile, the United Kingdom, and the United States of America. E.L.M.’s research is supported by The Spanish Ministry of Science via project AYA2007-67458.

REFERENCES

- Burgasser, A. 2007, *ApJ*, **658**, 617
- Burgasser, A., Geballe, T. R., Leggett, S. K., Kirkpatrick, J. D., & Golimowski, D. A. 2006, *ApJ*, **637**, 1067
- Burgasser, A., Tinney, C. G., Cushing, M. C., Saumon, D., Marley, M. S., Bennett, C. S., & Kirkpatrick, J. D. 2008, *ApJ*, **689**, L53
- Burgasser, A., et al. 2000, *ApJ*, **531**, L57
- Burgasser, A., et al. 2002, *ApJ*, **564**, 421
- Burningham, B., et al. 2008, *MNRAS*, **391**, 320
- Burrows, A., Sudarsky, D., & Lunine, J. I. 2003, *ApJ*, **596**, 587
- Cushing, M. C., et al. 2006, *ApJ*, **648**, 614
- Cushing, M. C., et al. 2008, *ApJ*, **678**, 1372
- Delorme, P., et al. 2008a, *A&A*, **482**, 961
- Delorme, P., et al. 2008b, *A&A*, **484**, 469
- Epchtein, N. 1998, in Proc. 179th Symp. Int. Astronomical Union (Dordrecht: Kluwer), 106
- Fazio, G., et al. 2004, *ApJS*, **154**, 10

- Fegley, B., & Lodders, K. 1996, [ApJL](#), **472**, L37
- Freedman, R. S., Marley, M. S., & Lodders, K. 2008, [ApJS](#), **174**, 504
- Golimowski, D. A., et al. 2004, [AJ](#), **127**, 3516
- Houck, J. R., et al. 2004, [ApJS](#), **154**, 18
- Hubeny, I., & Burrows, A. 2007, [ApJ](#), **669**, 1248
- Kirkpatrick, D. J., Beichman, C. A., & Skrutskie, M. F. 1997, [ApJ](#), **476**, 311
- Kirkpatrick, D. J., et al. 1999, [ApJ](#), **519**, 802
- Knapp, G. R., et al. 2004, [AJ](#), **127**, 3553
- Lawrence, A., et al. 2007, [MNRAS](#), **379**, 1599
- Leggett, S. K., Marley, M. S., Freedman, R., Saumon, D., Liu, M. C., Geballe, T. R., Golimowski, D. A., & Stephens, D. 2007a, [ApJ](#), **667**, 537
- Leggett, S. K., Saumon, D., Marley, M. S., Geballe, T. R., Golimowski, D. A., Stephens, D., & Fan, X. 2007b, [ApJ](#), **655**, 1079
- Leggett, S. K., et al. 2000, [ApJ](#), **536**, L35
- Liu, M. C., Leggett, S. K., & Chiu, K. 2007, [ApJ](#), **660**, 1507
- Luhman, K. L., et al. 2007, [ApJ](#), **654**, 570
- Marley, M. S., Seager, S., Saumon, D., Lodders, K., Ackerman, A. S., Freedman, R. S., & Fan, X. 2002, [ApJ](#), **568**, 335
- Martín, E. L., Delfosse, X., Basri, G., Goldman, B., Forveille, T., & Zapatero Osorio, M. 1999, [AJ](#), **118**, 2466
- McLean, I., et al. 2003, [ApJ](#), **596**, 561
- Mugrauer, M., Seifahrt, A., Neuhauser, R., & Mazeh, T. 2006, [MNRAS](#), **373**, L31
- Patten, B. M., et al. 2006, [ApJ](#), **651**, 502
- Saumon, D., & Marley, M. S. 2008, [ApJ](#), **689**, 1327
- Saumon, D., Marley, M. S., Cushing, M. C., Leggett, S. K., Roellig, T. L., Lodders, K., & Freedman, R. S. 2006, [ApJ](#), **647**, 552
- Saumon, D., et al. 2007, [ApJ](#), **662**, 1245
- Skrutskie, M. F., et al. 2006, [AJ](#), **131**, 1163
- Stephens, D. S., Leggett, S. K., Cushing, M. C., Saumon, D., Marley, M. S., Geballe, T., Golimowski, D. A., & Fan, X. 2009, [ApJ](#), submitted
- Strauss, M. A., et al. 1999, [ApJ](#), **522**, 61
- Tinney, C. G., Burgasser, A. J., Kirkpatrick, J. D., & McElwain, M. W. 2005, [AJ](#), **130**, 2326
- Tokunaga, A. T., Simons, D. A., & Vacca, W. D. 2002, [PASP](#), **114**, 180
- Toon, O. B., McKay, C. P., Ackerman, T. P., & Santhanam, K. 1989, [J. Geophys. Res.](#), **94**, 16287
- Vrba, F. J., et al. 2004, [AJ](#), **127**, 2948
- Warren, S. J., et al. 2007, [MNRAS](#), **381**, 1400
- Werner, M. W., et al. 2004, [ApJS](#), **154**, 1
- York, D. G., et al. 2000, [AJ](#), **120**, 1579